

Chemical Abundances and Doppler Imaging of the He-weak Ap Si+ Star HD 100357

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ABSTRACT

We present the results of time-resolved photometry, abundance analysis and Doppler Imaging of an Ap star, HD 100357. The *TESS* photometry revealed rotational modulation with a period of 1.6279247 days. Upon inspecting the residuals after removing the rotational period and its harmonics, we found additional frequencies around 15.8054 d^{-1} which we later confirmed with ground-based observations as originating from a nearby star. Using high-resolution spectroscopy, we identified HD 100357 as a He-weak Ap Si+ star exhibiting rotational modulation caused by surface abundance spots. The stellar parameters of HD 100357 were determined as $T_{\text{eff}} = 11,850 \text{ K}$, $\log g = 4.57$, $v \sin i = 60 \text{ km s}^{-1}$ and an inclination angle $i = 72^\circ$. The detailed abundance analysis revealed strongly overabundant stratified silicon, an overabundance of iron-peak elements and rare earth elements combined with remarkably deficient helium. Mapping of Fe and Cr abundances revealed the existence of ring-shaped regions with a lower concentration of the elements. Their geometry might reflect the orientation of the hypothetical magnetic field of the star, oriented $\sim 90^\circ$ to the rotational axis. HD 100357, with its strong chemical peculiarities and indications of high magnetic fields, represents an interesting candidate for follow-up spectropolarimetric observations aimed at investigating its magnetic field topology and stellar activities.

Key words: stars: abundances – stars: chemically peculiar – (stars:) starspots – line: profiles – stars: individual: HD100357 – stars: rotation

1 INTRODUCTION

Approximately 10% of main-sequence B- and A-type stars exhibit strong, globally organised magnetic fields (up to $\sim 30 \text{ kG}$) and unusually slow rotation ($v \sin i < 100 \text{ km s}^{-1}$) (Sikora et al. 2019), while some exceptions have faster rotations (Mikulášek et al. 2022). These stars, known as magnetic chemically peculiar (CP) stars, display surface abundance anomalies that correlate with effective temperature. The cooler CP stars, classified as Ap SrCrEu stars ($7000 \text{ K} \leq T_{\text{eff}} \leq 10\,000 \text{ K}$), show pronounced overabundances of Sr, Cr, and

rare earth elements, whereas the hotter Ap Si and He-weak stars ($10\,000 \leq T_{\text{eff}} \leq 15\,000 \text{ K}$) exhibit significant Si enhancements and He deficiencies (Landstreet et al. 2007; Romanyuk et al. 2013).

Preston (1974) divided CP stars into four main categories: Am (CP1), Ap (CP2), HgMn (CP3), and He-weak (CP4). Of these, CP2 and CP4 stars possess strong global magnetic fields, with strengths typically ranging from a few hundred gauss to several tens of kilogauss (Aurière et al. 2007; Elkin et al. 2010; Shultz et al. 2019). In contrast, CP1 and CP3 stars are of non-magnetic nature. He-weak stars, in particular, are typically early B-type stars that exhibit strong underabundances of helium and overabundances of elements such as Si, P, and Ga (Norris 1971; Zboril & North 1997). The He deficiency

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in these stars is believed to arise from helium gravitational settling and inhibited mixing in the presence of magnetic fields (Vauclair & Vauclair 1975).

The anomalies in CP stars are explained by atomic diffusion, in which radiative levitation of certain heavy ions and gravitational settling of lighter ions occur in the absence of convective mixing (Michaud 1970; Richard et al. 2001; Alecian 2018a,b). The quenching of mixing in magnetic stars are thought to result from magnetic braking that slows stellar rotation during early evolutionary stages (Fleck 1980), promoting atmospheric stability.

The magnetic fields in CP stars not only govern diffusion but also suppress differential rotation and convective turbulence, producing horizontal abundance patches across the stellar surface. These spots rotate with the star, giving rise to photometric and spectroscopic variability on rotational timescales. Such variability can be used to infer stellar rotation periods and map surface inhomogeneities via Doppler Imaging (DI) and Zeeman Doppler Imaging (ZDI; Kochukhov et al. 2004; Donati et al. 2006). Rotation periods of magnetic CP stars span a wide range, from less than a day to several decades, with very long periods associated with the most strongly magnetic and slowly rotating examples (Mathys 2017; Giarrusso et al. 2022; Mathys et al. 2024).

In addition to surface inhomogeneities, vertical stratification in the stellar atmosphere causes different chemical elements to concentrate at different depths. This enables the use of lines forming at different atmospheric layers to probe depth-dependent pulsation modes, particularly in a cooler subset of the magnetic CP stars that pulsate in high-overtone, low-degree, non-radial p -modes aligned with their magnetic axes, known as the rapidly oscillating Ap (roAp) stars (Baldry et al. 1998; Kochukhov & Ryabchikova 2001; Mkr-tichian et al. 2003; Ryabchikova et al. 2004). The strong chemical peculiarities and magnetic field in the presence of pulsations provide great insights into stellar interiors (Joshi & Joshi 2015; Kurtz 2022) and are among the best laboratories for studying magneto-acoustic interactions in stellar atmospheres (Kurtz 2000; Saio 2005).

The Nainital–Cape (N-C) Survey is one of the most extensive ground-based campaigns aimed at detecting and characterizing rapidly oscillating Ap (roAp) stars in the northern and southern hemispheres. Using the 1.04-m Sampurnanand Telescope at ARIES, Nainital, and the 0.5-m telescope at the South African Astronomical Observatory (SAAO), hundreds of targets were monitored over nearly two decades. The key results of this systematic survey are summarised in Ashoka et al. (2000); Martinez et al. (2001); Joshi et al. (2003, 2006, 2009, 2010, 2012, 2016, 2017, 2022) and have contributed significantly to the discovery of new roAp candidates and the understanding of the low-amplitude variability in CP stars. We continue this study by investigating the photometric variability in CP star candidates from the survey through space-based missions.

In this paper, we present a comprehensive investigation of the southern He-weak Ap Si+ star HD 100357 (reported a null result for pulsation in the N-C survey) using both ground-based photometric and spectroscopic observations, as well as space-based data from the *TESS* mission. This work expands the sample of He-weak Ap Si+ stars in the southern sky, which serve as a transitional class between CP2 and CP4 stars, providing insights into the interactions among magnetism, rotation, and chemical stratification in CP stars.

The manuscript is structured as follows. Sec. 2 discusses the observations and data preparation. In Sec. 3, we determine the stellar parameters for the star. The chemical compositions and determination of the surface distribution of Fe and Cr are outlined in Sec. 4 and Sec. 5, respectively. Finally, the summary and conclusion drawn from our findings are given in Sec. 6.

Table 1. Observational log of *TESS* data. The cadence and duration of observations are given in seconds (s) and days (d), respectively.

TESS Sector	Observation Period	Cadence (s)	Duration (d)
10	March 26 – April 21, 2019	120	26
11	April 22 – May 20, 2019	1800	28
37	April 2 – April 28, 2021	600	26
38	April 29 – May 26, 2021	600	27
64	April 6 – May 4, 2023	120	28
65	May 4 – June 1, 2023	120	28

2 OBSERVATIONS AND DATA ANALYSIS

2.1 Photometry

2.1.1 *TESS*

TESS is an all-sky survey telescope whose primary objective is to detect exoplanets orbiting nearby bright stars using the transit method (Ricker et al. 2015). In addition to the primary mission, *TESS* also provides high-precision time-series photometry suitable for studying other types of variability, such as rotational, pulsational, or eclipses of stars brighter than 12^{th} magnitude (Sullivan et al. 2015; Bowman et al. 2022). *TESS* Target Pixel Files (TPFs) and Light Curve Files are available in short cadences of 20 and 120 s, while Full-Frame Images (FFIs) are obtained in long cadences: 1800 s for Sectors 1-26 (Cycles 1-2), 600 s for Sectors 27-55 (Cycles 3-4), and 200 s from Sector 56 onwards (Cycle 5+). These data products are publicly available and can be downloaded through the Mikulski Archive for Space Telescopes (MAST)¹.

HD 100357 was observed in six different sectors by *TESS* and the log of observations is given in Table 1. The *TESS* light curve files contain two types of flux, namely, the Simple Aperture Photometry (SAP) flux and the Pre-search Data Conditioning SAP (PDCSAP) flux. In the PDCSAP flux, long-term instrumental trends are removed using Co-trending Basis Vectors (CBVs) (Smith et al. 2012). For the present study, we used 120 s, 600 s, and 1800 s cadence PDCSAP flux downloaded from the MAST archive via the `LIGHTCURVE` Python package (Lightkurve Collaboration et al. 2018).

The light curve fluxes were converted to relative magnitudes (mmag) and centred at zero by subtracting the mean magnitude. Light curves were then stitched together and subjected to frequency analysis to identify the rotational frequency using the discrete Fourier fitting technique implemented in the `PERIOD04` package (Lenz & Breger 2005, 2014). We fitted the fundamental rotational frequency and its 15 consecutive integer multiples (harmonics) simultaneously using a multi-component sinusoidal model and performed subsequent prewhitening to identify any higher frequencies. A signal-to-noise ratio (S/N) greater than 5.3 was adopted as the criterion for selecting significant frequencies, based on the detection threshold for *TESS* sectors by Baran & Koen (2021). The derived frequencies, amplitudes, and S/N values obtained are listed in Table 2. The light curve, frequency spectrum and residuals are shown in Fig. 1.

HD 100357 (TIC 280667311) was classified as an Ap star with A0 spectral type and showing EuCrSr type peculiarity by Houk & Cowley (1975). These stars are known to show rotational variability due to chemical spots. Bernhard et al. (2020) reported a rotational

¹ <https://archive.stsci.edu/>

Table 2. Frequency, amplitude and S/N from the combined light curves of all six sectors. f_{rot} is the rotational frequency, and the harmonics are given as integer multiples of this frequency. The labels ‘ f_{n*} ’, $n = 1, 2$ and 3 indicate the frequencies of the contaminating star.

	Frequency (d^{-1})	Amplitude (mmag)	S/N
$2f_{\text{rot}}$	1.2285581(3)	17.733(6)	545.2
f_{rot}	0.6142790(1)	8.520(6)	140.8
$4f_{\text{rot}}$	2.457116(3)	0.762(6)	54.8
$3f_{\text{rot}}$	1.842837(8)	0.281(6)	15.0
$5f_{\text{rot}}$	3.07140(2)	0.122(6)	10.1
$7f_{\text{rot}}$	4.29995(2)	0.099(6)	11.7
$10f_{\text{rot}}$	6.14279(3)	0.085(6)	11.7
$12f_{\text{rot}}$	7.37135(4)	0.060(6)	9.1
$8f_{\text{rot}}$	4.91423(4)	0.059(6)	7.3
$14f_{\text{rot}}$	8.59991(4)	0.052(6)	8.9
f_{1*}	15.80501(3)	0.087(6)	15.6
f_{2*}	14.32267(6)	0.038(6)	6.5
f_{3*}	16.98574(7)	0.034(6)	5.7

period of 1.62794 d for HD 100357 star using ground-based ASAS-3 data. In our frequency analysis, we detected a rotational period in agreement with the previous value and nine additional harmonics. Interestingly, the amplitude of the first harmonic is higher than the fundamental rotational frequency, indicating the presence of multiple spots located on diametrically opposite sides of the star surface with complex geometry. We performed a parabolic fit to the peak of each light maximum identified in all *TESS* sectors to obtain the accurate timings and finally fitted the linear ephemeris for this star. The updated linear ephemeris is:

$$\text{BJD}_{\text{max}} = 2458570.9785(3) + 1.6279247(7) \times E, \quad (1)$$

where BJD_{max} is the time of light maxima at epoch E .

2.1.2 TRAPPIST-South

We conducted ground-based follow-up observations in the field surrounding HD 100357 using the TRAPPIST-South telescope. It is a 60-cm Ritchey-Chretien telescope located at ESO-La Silla Observatory in Chile. It is equipped with a thermoelectrically cooled 2K×2K FLI Proline CCD camera with a field of view of $22' \times 22'$ and a pixel scale of $0.65''/\text{pixel}$ (Jehin et al. 2011; Gillon et al. 2011). The observations were carried out on 18 March 2025 over a duration of four hours, with individual exposures of 10 s, through the standard Johnson–Cousins R -band filter. We used the python package Prose² (Garcia et al. 2022) for analyzing the images and subsequently performing differential photometry. The primary aim of these observations was to assess the possible contamination in the *TESS* light curve of HD 100357.

2.1.3 Contamination

After fitting and subtracting the rotational frequency and its harmonics from the time-series data, the residuals show low-amplitude frequencies at 15.805, 14.323, and 16.986 d^{-1} (see inset of middle

² <https://prose.readthedocs.io/en/latest/ipynb/sources.html>

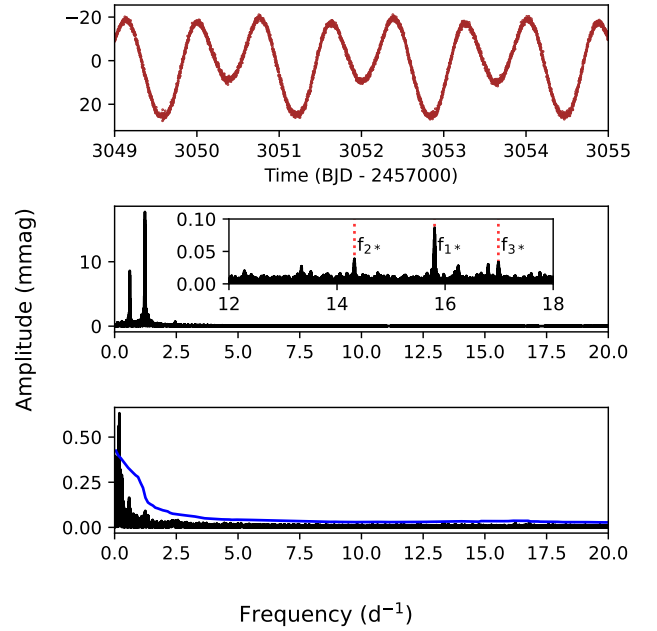


Figure 1. The top panel displays a segment of the *TESS* light curve with a cadence of 120 seconds. The second panel shows the amplitude spectrum using the combined light curve from all *TESS* sectors. The inset in the second panel shows the frequency range from 12 to 18 d^{-1} , where contaminating signals in the residuals persist after fitting and eliminating 15 successive harmonics of the fundamental rotational frequency. The third panel shows the residual after removing all the significant frequencies, where the blue line represents the 5.3 S/N level. The low-frequency peaks in the residuals can be instrumental artefacts.

panel of Fig. 1). Since the *TESS* pixels have poor spatial resolution (21 arcsec), the contamination of signals from nearby bright stars is expected. To mitigate this, we performed a pixel-by-pixel investigation of the *TESS* target pixel file to check for any contamination. We used the technique described in Dileep et al. (2025), by selecting custom apertures around the target star identified through Gaia DR3 positions (Gaia Collaboration et al. 2023) (See Fig. B1). Our finding is that the additional frequencies detected in the residuals of HD 100357 ($G_{\text{mag}} = 8.98$) originated from a nearby star (Gaia DR3 5236626819707864192; $G_{\text{mag}} = 13.01$; $\Delta G_{\text{mag}} = 4.03$), possibly a new δ Scuti star. We further confirmed the origin of these frequencies with ground-based observations from the TRAPPIST-South telescope, which has a spatial resolution 30 times better than that of *TESS*, allowing the contaminating star to be fully resolved in the corresponding CCD frames. Fig. B2 shows the folded light curve using the dominant contaminating frequency of 15.8054 d^{-1} , along with a CCD image of the field that marks the stars. Apart from this, there is a very close visual companion to the star that can cause contamination, which will be discussed in the following section.

2.2 GAIA Astrometry

HD 100357 was classified as a visual double star in the Washington Double Star Catalogue (WDS) (Mason et al. 2001) and the Tycho Double Star Catalogue (Fabricius et al. 2002). Gaia DR2 (Gaia Collaboration et al. 2018) and DR3 (Gaia Collaboration et al. 2023) catalogues also reported this companion star. From the Gaia DR3 position and parallax, they are separated by ~ 2.55 arcsecs in the sky and ~ 5.33 pc apart in space. Kervella et al. (2019) and Brandt

(2021) searched for proper motion anomalies in several stars using the Hipparcos and Gaia proper motions and parallax, but they did not report any significant proper motion anomaly for HD 100357. Hence, we can conclude that these are comoving stars with no orbital interaction. Nevertheless, the nearby visual companion with a *Gaia* G magnitude of 12.35 compared to the target's $G = 8.98$, may cause flux contamination in the *TESS* pixels at a level of approximately 5.7%. Since we did not detect any additional variability from this nearby companion, it will only add a constant noise level to the signal from HD 100357; hence, one can neglect the contamination in this case.

2.3 Spectroscopy

We acquired 20 high-resolution spectra for the spectroscopic analysis of HD 100357. One spectrum was taken with the High Resolution Échelle Spectrograph (HRS) (Tyas 2012) attached to the 10.0-m Southern African Large Telescope (SALT) of the South African Astronomical Observatory (Sutherland, South Africa) with spectral resolution $R \approx 37,000$ over the wavelength range 370–890 nm. We retrieved one archival spectrum with a spectral resolution $R \approx 48,000$ over 350–920 nm, from the Science Portal of the European Southern Observatory (ESO)³ that was observed with the Fiberfed Extended Range Optical Spectrograph (FEROS) (Kaufer et al. 1999) at MPG/ESO 2.2-m telescope (La Silla, Chile). Finally, a dedicated program was carried out with the High Efficiency and Resolution Canterbury University Large Échelle Spectrograph (HERCULES) (Hearnshaw et al. 2002), equipped on the 1-m McLellan telescope of Mount John University Observatory (Mount John, New Zealand). In total, 18 HERCULES spectra were collected with a spectral resolution of $R \approx 41,000$ over the wavelength range 380–880 nm. The standard reduction, wavelength calibration, and continuum normalisation of HERCULES spectra were performed using the dedicated pipeline HRSP (Skuljan 2004), while the HRS spectrum was reduced and wavelength calibrated using the SALT HRS pipeline (Crawford et al. 2010; Kniazev et al. 2016, 2017). The HRS and FEROS spectra were then continuum normalised using the CONTINUUM task of the Image Reduction and Analysis Facility⁴ (IRAF). For the in-depth spectroscopic analysis, all these spectra were corrected for the barycentric motion. We calculated the S/N for each spectrum at a wavelength of 5500 Å. The log of spectroscopic observations is given in Table 3.

2.3.1 Least-squares Deconvolution

Magnetic CP stars show significant line profile variations as a result of chemical spots on their surface. These variations are difficult to detect when the S/N of a spectrum is very low. To overcome this difficulty, the Least-Squares deconvolution (LSD) technique (Donati & Brown 1997; Kochukhov et al. 2010) is used to get a higher S/N mean line profile using a large number of metallic lines. We constructed the LSD profiles for the observed spectra of HD 100357 using a line mask prepared with the atomic data extracted from the VALD3 database (Piskunov et al. 1995; Ryabchikova et al. 2015; Pakhomov et al. 2019). We avoided the broad Balmer lines and the interstellar Na lines for constructing the LSD profiles. We fitted a rotationally broadened profile to the LSD profiles, excluding the core areas (see Fig. 2), where line profile variations occur, to derive the radial velocities and $v \sin i$ values for each observation. The results

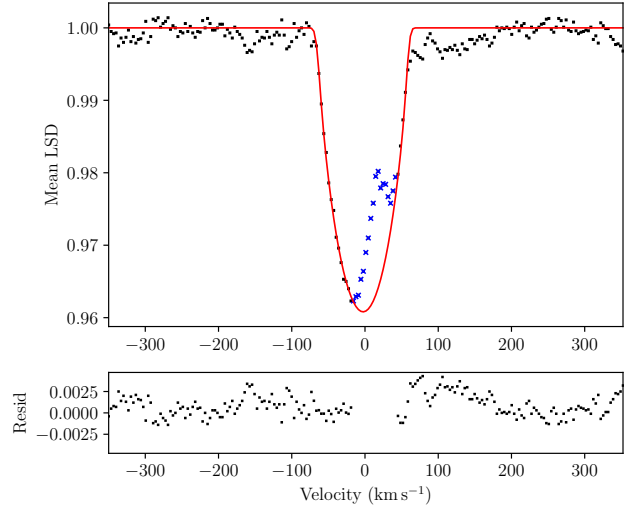


Figure 2. Top panel: The scattered points represent the LSD profile of HD 100357 obtained from SALT spectrum. The solid curve is the best-fit rotationally broadened line profile. The points in the middle, represented by cross symbols, are excluded from the line profile fit. Bottom Panel: The residuals of the best-fit line profile to the observed LSD profile.

derived from the LSD profiles are shown in the last two columns of Table 3.

3 STELLAR PARAMETERS

3.1 Photometric Parameters

To estimate the effective temperature, T_{eff} , and surface gravity, $\log g$, of our target, we used photometric data published in the General Catalogue of Photometric Data⁵ (Mermilliod et al. 1997) for the Strömgren and Geneva systems. HD 100357 is located close to the Galactic equator, a region having significant interstellar reddening. A colour excess of $E(B - V) = 0.10$ mag was determined from the intensity of the Na D1 line, following the calibration by Munari & Zwitter (1997), and the parallax is $\pi = 2.72 \pm 0.02$ mas (Gaia Collaboration et al. 2021). In addition, we used the 3D dust map G-Tomo⁶ (Lallement et al. 2022; Vergely et al. 2022) to derive the colour excess $E(B - V) = 0.136$, which is in good agreement with the value derived from the Na D1 line.

Using the calibrations given by Napiwotzki et al. (1993) for indices in Strömgren system results in the estimated $T_{uvby} = 12,420$ K and $\log g = 4.56$. At the same time, the effective temperature calibrated using index $[u - b]$ turns out to be 1,000 K lower: 11050 K. It is the typical feature of CP stars with strong chemical anomalies. In light of anomalies, it sounds reasonable to apply the special relation from Eq. (12) of the original publication. This brings us $T_{[u-b]}(\text{Ap}) = 10750$ K. Unfortunately, Napiwotzki et al. (1993) did not provide the method to estimate the errors. Following the common practice, it would be reasonable to assume that the typical error of T_{eff} is close to 300 K, while for $\log g$ one can expect an error of order 0.15. To account for the interstellar reddening in the Geneva data, we converted $E(B - V)$ to $E(B2 - V1)$ by dividing it by 1.14.

³ <https://archive.eso.org/scienceportal/home>

⁴ <https://iraf-community.github.io/>

⁵ <https://gcpd.physics.muni.cz/>

⁶ <https://explore-platform.eu/sdas/about/gtomo>

Table 3. Log of spectroscopic observations of HD 100357 in the chronological orders. The rotational phase, radial velocity (RV) and projected rotational velocity ($v \sin i$) are also listed. The rotational phases were calculated with reference time, $t_0 = 2458570.9785$ BJD.

Telescope	Diameter (m)	Spectrograph	Phase	Date	BJD (2400000+)	Resolution	Integration (s)	S/N (@550nm)	RV (km s ⁻¹)	vsini (± 2 km s ⁻¹)
MPG/ESO	2.2	FEROS	0.938	08 Feb 2007	54139.6498	48000	800	133	-4.2 ± 0.3	56
SALT	10.0	HRS	0.334	14 Jan 2023	59958.5137	37000	300	74	-3.1 ± 0.4	63
McLellan	1.0	HERCULES	0.614	25 Dec 2023	60304.0915	41000	1800	45	2.7 ± 0.5	56
			0.628	25 Dec 2023	60304.1135	41000	1800	51	-1.5 ± 0.5	60
			0.839	27 Dec 2023	60306.0830	41000	1800	18	-2.0 ± 1.0	63
			0.235	31 Dec 2023	60309.9862	41000	1800	41	-1.6 ± 0.6	62
			0.249	31 Dec 2023	60310.0082	41000	1800	34	-1.2 ± 0.5	61
			0.264	31 Dec 2023	60310.0324	41000	1800	33	-2.1 ± 0.5	63
			0.277	31 Dec 2023	60310.0544	41000	1800	45	-1.7 ± 0.4	63
			0.883	01 Jan 2024	60311.0397	41000	1800	53	-1.6 ± 0.4	61
			0.896	01 Jan 2024	60311.0617	41000	1800	53	-1.1 ± 0.4	61
			0.513	02 Jan 2024	60312.0662	41000	1800	59	-1.6 ± 0.4	60
			0.528	02 Jan 2024	60312.0896	41000	1800	44	-0.9 ± 0.4	60
			0.461	23 Jan 2024	60333.1452	41000	1800	40	-4.0 ± 0.5	59
			0.925	19 Feb 2024	60359.9473	41000	1800	30	-3.0 ± 0.6	58
			0.939	19 Feb 2024	60359.9693	41000	1800	39	-3.3 ± 0.5	57
			0.561	20 Feb 2024	60360.9817	41000	1800	67	0.1 ± 0.4	60
			0.574	20 Feb 2024	60361.0037	41000	1800	63	0.9 ± 0.4	58
0.184	21 Feb 2024	60361.9966	41000	1800	65	-1.6 ± 0.4	61			
0.198	21 Feb 2024	60362.0186	41000	1800	55	-1.5 ± 0.4	61			

Calibrations published by [Kunzli et al. \(1997\)](#) for the Geneva system, depending on the expected metallicity, give two sets of data: $T_{\text{Geneva}} = 11,726$ K and $\log g = 4.59$ for the case of $[M/H] = 0$, or $T_{\text{Geneva}} = 11,560$ K and $\log g = 4.56$ when $[M/H] = +1$. The estimated error of determination is 75 K for T_{eff} and 0.07 for $\log g$.

As the problem of finding basic parameters from photometry is especially acute, we also revised the mentioned parameters found from the original calibrations in the framework of the study by [Netopil et al. \(2008\)](#). This gives the corrected values: $T'_{\text{uvby}} = 11,480$ K and $T'_{\text{Geneva}} = 10,975$ K.

To summarise, we can state that photometry cannot give consistent estimations of the effective temperature. We can sort all results into two groups with the mean values: $T_{\text{orig}} = 11,900$ K and $T_{\text{Ap}} = 11,070$ K. The first value is used as the effective temperature of HD 100357 through this study. However, we cannot exclude that the real T_{eff} can be lower by about 1,000 K. As for the surface gravity, we adopted the mean value, $\log g = 4.57 \pm 0.15$, which still appears overestimated.

3.2 Distance, Luminosity, and Radius

The distance of a star can be determined by the Gaia DR3 parallax. The parallax angle $\pi = 2.72 \pm 0.02$ mas results in the distance of our target $d = 367 \pm 3$ pc. We calculated the extinction in the V band $A_V = R_V \times E(B - V)$, where $R_V = 3.1$ ([Cardelli et al. 1989](#)). We estimated the bolometric correction (BC = -0.65) using the empirical relation given by [Torres \(2010\)](#). For calculating the Johnson V magnitude of the star, we used the conversion from Tycho V_T and B_T magnitude ([Bessell 2000](#)), resulting in the apparent magnitude, $V = 8.97 \pm 0.02$. This yields the value of $\log(L/L_{\odot}) = 1.86 \pm 0.02 L_{\odot}$. The radius of the star is calculated from the Stefan-Boltzmann equation, and the value is $R = 2.03 \pm 0.10 R_{\odot}$.

3.3 Inclination Angle

The inclination angle of the rotation axis of a magnetic CP star is required to model the abundance spot on the surface of the star and determine the geometry of the magnetic fields ([Stibbs 1950](#); [Kochukhov 2017](#)). The equatorial velocity of a rotating star can be calculated as $v_{\text{eq}} = 2\pi R/P$, where R is the radius of the star and P is its rotational period. We used the rotational period determined from the *TESS* data, and mean $v \sin i$ (60 ± 2 km s⁻¹) calculated from the spectroscopy, giving an inclination angle, $i = 72^\circ \pm 11^\circ$.

3.4 Spectroscopic Parameters

The absorption-line spectrum of HD 100357 aligns with that of an early A-type photosphere. Comparison with the spectral standards from [Gray & Corbally \(2009\)](#) shows that He I lines are weak across a wide spectral range. In contrast, lines of Si II, Sc II, Ti II, Cr II, Mn II, Fe II, Co II, and Sr II are unusually strong. Notably, the Mg II 4481 Å line, which is typically one of the strongest, appears relatively shallow in the spectrum of HD 100357, which is a typical feature of He-weak stars.

The atmospheric parameters of HD 100357 were determined using a self-consistent iterative approach based on spectral synthesis, utilising the SALT HRS spectra. As a starting point, we adopted the effective temperature ($T_{\text{eff}} = 11,900$ K) and surface gravity ($\log g = 4.57$) derived from photometric data. In the subsequent step, a He-weak atmospheric model with a hydrogen abundance of $N_{\text{H}}/N_{\text{tot}} = 0.99$ was computed using the ATLAS12 model atmosphere code ([Castelli 2005](#); [Kurucz 2005](#)) that accounts for the impact of chemical composition anomalies on the atmospheric opacity distribution.

Theoretical spectra were computed assuming local thermodynamic equilibrium (LTE) using this model atmosphere and atomic data obtained from the VALD3 database ([Piskunov et al. 1995](#); [Ryabchikova et al. 2015](#); [Pakhomov et al. 2019](#)) with the

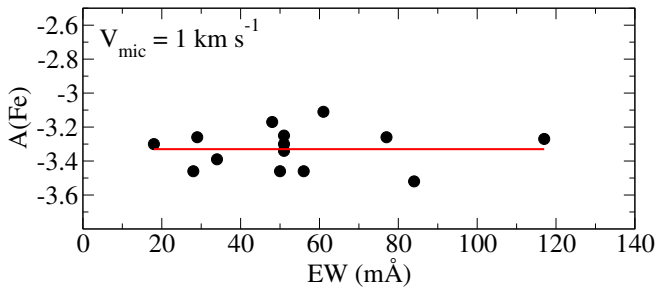


Figure 3. The scatter points represent iron abundances as a function of their equivalent widths for HD 100357. The solid straight line is the optimal minimised slope of this relationship at $\xi_{\text{mic}} = 1 \text{ km s}^{-1}$.

synthV_NLTE code (Tsymbal et al. 2019). To compare the theoretical stellar spectra with observational data, we utilised the widget program BinMag⁷ (Kochukhov 2018) that interfaces to the synthV_NLTE code and allows for automatic determination of the best fits to the observed line profiles, enabling determination of highly accurate chemical abundances.

The T_{eff} and $\log g$ parameters were refined by comparing the synthetic spectrum to the observed spectrum in the wings of the H_{α} and H_{β} hydrogen lines. Fig. 4 presents a comparison between the observed and synthetic profiles calculated using this final set of parameters. The microturbulent velocity, $\xi_{\text{mic}} = 1 \text{ km s}^{-1}$, was determined using the classical method of minimising the slope in the relationship between elemental abundance and equivalent widths (see Fig. 3). The equivalent widths were derived using synthetic line-profiles, accounting for blending due to rotational broadening (60 km s^{-1}), ensuring that the contribution of each line was properly modeled.

Our analysis confirmed that the initial photometry-based parameters are consistent with the spectroscopic examination of the hydrogen lines. As a result, we determined the following parameters: effective temperature $T_{\text{eff}} = 11,850 \pm 150 \text{ K}$, surface gravity $\log g = 4.57 \pm 0.05$, and a microturbulent velocity $\xi_{\text{mic}} = 1.0 \pm 0.2 \text{ km s}^{-1}$. The iron abundance is measured using 14 lines, and the average value is $\log A(\text{Fe}) = -3.33 \pm 0.12$, which reflects an overabundance relative to solar with $[\text{Fe}/\text{H}] = 1.16 \text{ dex}$ (Lodders 2021).

3.5 Evolutionary Status

To determine the evolutionary status of the star, we constructed a Hertzsprung-Russell (H-R) diagram with theoretical isochrones and evolutionary tracks downloaded from MESA Isochrones and Stellar Tracks (MIST) (Dotter 2016; Choi et al. 2016) compiled using the Modules for Experiments in Stellar Astrophysics (MESA) code (Paxton et al. 2011, 2013, 2015, 2018). We placed our star in the H-R diagram using the determined values of $\log T_{\text{eff}}$ and $\log(L/L_{\odot})$ with associated errors. The position of HD 100357 in the HR diagram (Fig. 5) along with the inspection of its high-resolution spectrum in Sec. 3.4, indicates that it is a main-sequence star. From the inspection of the isochrones and stellar tracks, one can conclude that HD 100357 has a mass, $M \sim 2.83 M_{\odot}$ and an age, $t \sim 90 \text{ Myr}$. The stellar parameters determined from our analysis in this section are listed in Table 4.

Table 4. Summary of all stellar parameters of HD 100357 as ascertained from the current study. The associated error bars for the last digits of the relevant values are shown in brackets.

Parameters	values
P_{rot} (days)	1.6279294 (7)
t_0 (BJD)	2458570.9743 (4)
Vmag	8.97 (2)
E(B-V)	0.136 (2)
$\log(L/L_{\odot})$	1.86 (2)
T_{eff} (K)	11850 (150)
$\log g$	4.57 (5)
[Fe/H]	1.16 (12)
$v \sin i$ (km s^{-1})	60 (2)
ξ_{mic} (km s^{-1})	1.0 (2)
R (R_{\odot})	2.03 (10)
i (deg)	72 (11)
M (M_{\odot})	2.83 (4)
Age (Gyr)	0.09 (4)

4 CHEMICAL COMPOSITION

We used the SALT HRS spectrum to determine the chemical abundances of HD 100357, which was obtained in the most recent epoch and with the best S/N. Elemental abundances were generally derived from multiple absorption lines within the observed spectral range. Abundances are expressed as $\log(A)_X = \log(N_X / N_H)$, where N_X denotes the number density of element X and N_H represents the number density of hydrogen. The abundances were determined under the LTE assumption by fitting synthetic line profiles to observed ones using the BinMag tool with the line list extracted from the VALD3 database. Hyperfine splitting (Pakhomov et al. 2019) was also incorporated during the extraction of the line list. The mean abundances of each element X were calculated by averaging many lines. Uncertainties were determined as the standard deviation of the mean value when more than two lines were available. Relatively high rotational broadening complicates the task of selecting lines. Consequently, most elements in the HD 100357 spectrum are represented by a limited number of blended lines.

We derived the LTE abundances for 19 elements. The results of this analysis are summarized in Table 5 and illustrated in Fig. 6, which compares the derived abundances to solar abundances using the notation $[X/\text{H}] = \log(N_X/N_H) - \log(N_X/N_H)_{\odot}$. The solar reference abundances are taken from Lodders (2021). The detailed abundances from individual lines are listed in Table A1.

4.1 Helium

We identified weak absorption lines of He I at 4471 \AA and 5876 \AA and assessed the helium abundance as $\log(A)_{\text{He}} = -3.1 \text{ dex}$, which is 2 dex lower than the solar value. Fig. 7 presents a comparison between the observed and synthetic He I profiles computed using three distinct sets of He abundances for comparison.

4.2 Carbon, Oxygen, Magnesium, and Aluminium

We were able to measure only one C II line at 4267.259 \AA and the oxygen triplet in the wavelength range $7771\text{--}7775 \text{ \AA}$. Carbon and oxygen are both strongly depleted by more than 1.0 dex in the atmosphere of HD 100357. Slightly less depletion was found for Mg and Al, by $\sim 0.4 \text{ dex}$ compared to solar values.

⁷ <http://www.astro.uu.se/~oleg/binmag.html>

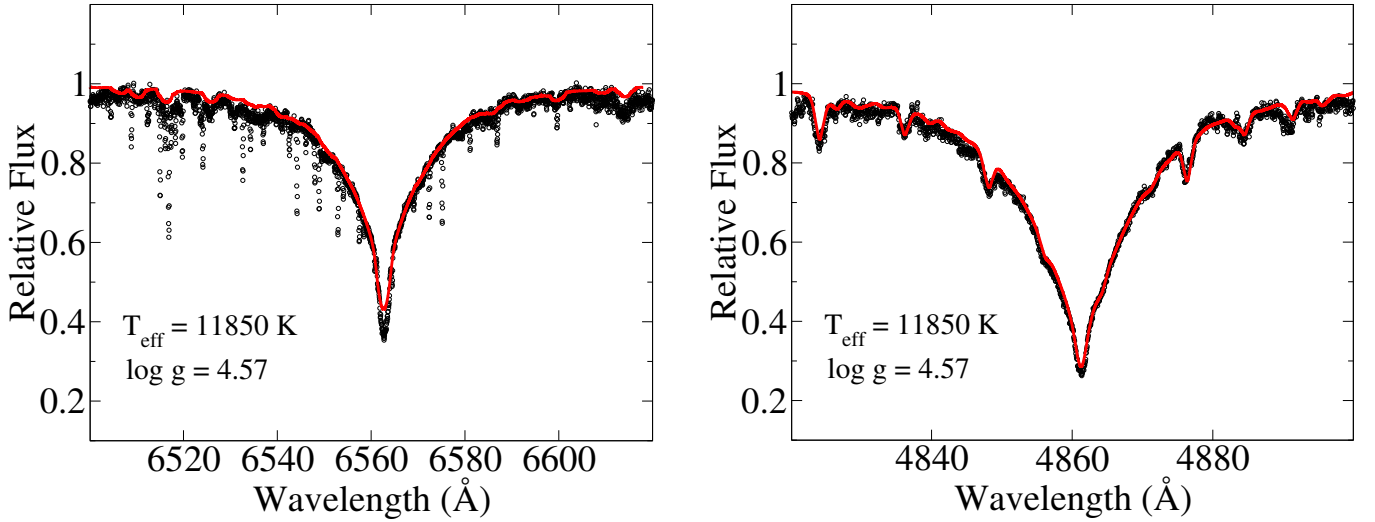


Figure 4. Comparison of the observed spectrum (black circles) and synthetic spectrum (red line) in the region of H_{α} (left panel) and H_{β} (right panel). The telluric lines near the H_{α} region were excluded during the fitting.

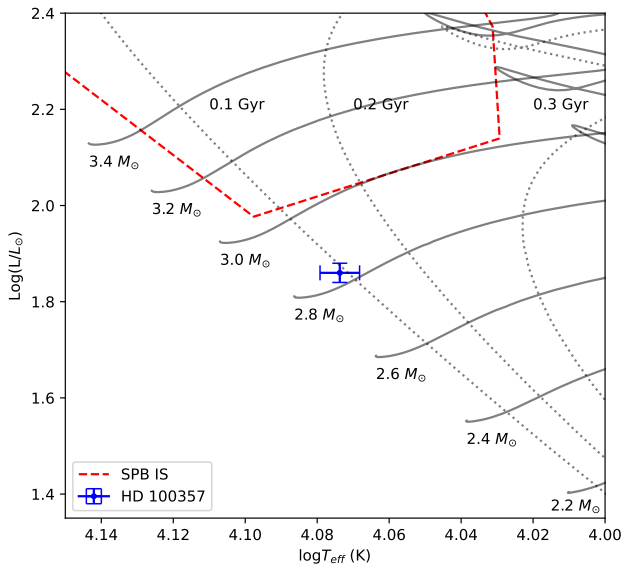


Figure 5. Position of HD 100357 in the H-R diagram with error bars. The dotted lines and the solid lines represent the isochrones and the evolutionary tracks, respectively. The dashed line is the instability strip of slowly pulsating B stars (Miglio et al. 2007).

4.3 Silicon

We were able to measure only one weak Si III line, which agrees well within the error bar with the mean abundance derived from Si II lines. Despite our best efforts to select primarily unblended Si II lines with varying excitation energies and oscillator strengths, we observed a significant scatter of 0.30 dex around the mean value. The big scatter in the Si II-based abundances prompted us to investigate the relationship between abundance measurements from individual lines and the line strength, which depends on oscillator strength and excitation energy (see, e.g., Ryabchikova (2014)). This relationship is illustrated in Fig. 8 for Si II lines. The figure shows a noticeable trend: stronger lines formed higher in the atmosphere yield lower

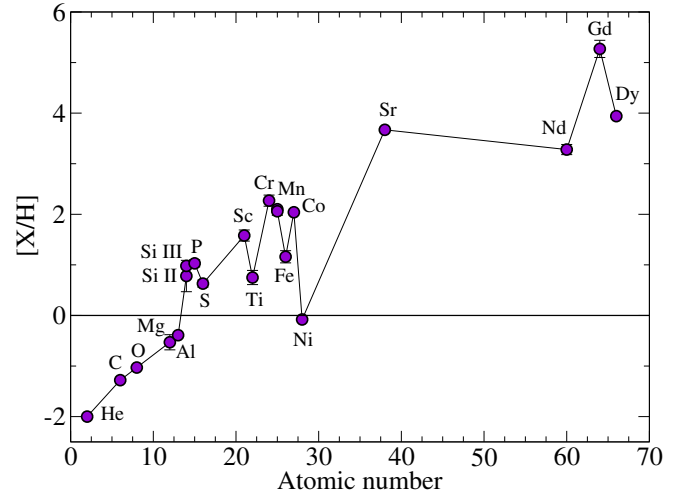


Figure 6. Relative photospheric abundance of chemical species of HD 100357 at rotational phase 0.334. The horizontal line at zero denotes the solar reference. There is an excess of Si, Fe-peak elements, and heavier elements, accompanied by a deficiency of light elements and a pronounced underabundance of He.

individual abundances compared to weaker lines formed closer to the photosphere. These results suggest a clear vertical stratification of Si, where most of the Si is concentrated in the deeper layers of the star (Bailey & Landstreet 2013). A detailed analysis of vertical abundance stratification is beyond the scope of this paper.

4.4 Phosphorous, Sulphur

We used two lines of P II at 4602 Å and 6166 Å and two lines of S II at 5454 Å, and 5606 Å. Our estimates indicate that both P and S are overabundant, with the abundance for P being 1.03 dex higher than the solar value and a slightly lesser abundance for S by 0.63 dex.

Among the two categories of He-weak stars — magnetic and non-magnetic — non-magnetic He-weak stars occasionally show overabundances of phosphorus (up to twice the solar value) and gallium

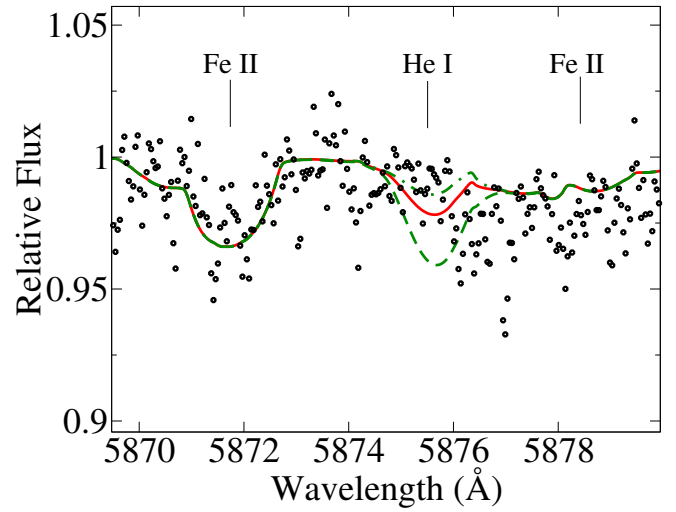
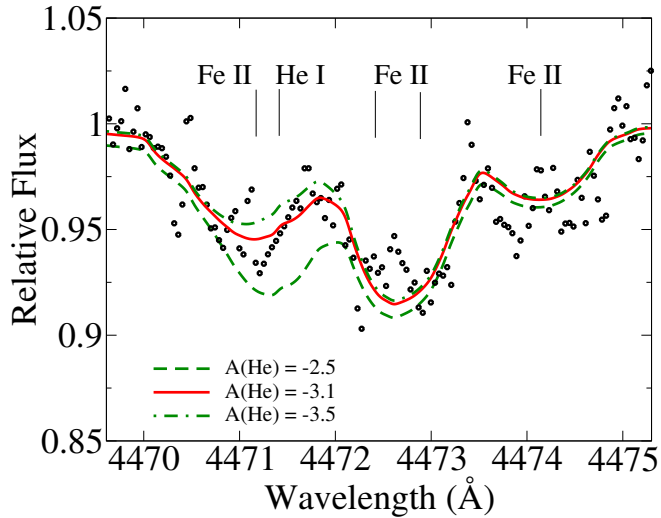


Figure 7. Observed He I $\lambda 4471 \text{ \AA}$ (left panel) and He I $\lambda 5876 \text{ \AA}$ (right panel) line profiles for HD 100357 (black circles), compared to synthetic profiles (red and green lines). The red solid lines were obtained assuming the best He abundance $A(\text{He}) = -3.1$. The green dashed lines show the synthetic profile by increasing He to $A(\text{He}) = -2.5$ value. The green dashed-dot lines show the synthetic profile by decreasing He to -3.5 . The nearby blended lines of Fe II are also indicated.

Table 5. The mean LTE abundances of chemical species in HD 100357. The parameter $[X/H] = \log(N_X/N_H) - \log(N_X/N_H)_\odot$, compares the observed abundances with solar abundances (Lodders 2021). The last column of the table lists the standard deviations for each species when more than two lines are available.

Atomic number	Ion	Number of lines	$\log(A)_X$	$[X/H]$	Sigma
2	He II	2	-3.1	-2	-
6	C II	1	-4.7	-1.28	-
8	O I	1	-4.27	-1.03	-
12	Mg II	3	-5.03	-0.53	0.15
13	Al II	3	-5.98	-0.39	0.03
14	Si II	19	-3.7	0.78	0.31
14	Si III	1	-3.5	0.98	-
15	P II	2	-5.55	1.03	0.07
16	S II	2	-4.25	0.63	0.07
21	Sc II	3	-7.27	1.58	0.11
22	Ti II	4	-6.28	0.75	0.14
24	Cr II	11	-4.11	2.27	0.11
25	Mn I	1	-4.45	2.1	-
25	Mn II	11	-4.49	2.06	0.08
26	Fe II	14	-3.33	1.16	0.12
27	Co II	2	-5	2.04	-
28	Ni I	3	-5.8	-0.08	0.01
38	Sr II	1	-5.41	3.67	-
60	Nd III	2	-7.27	3.28	0.1
64	Gd II	2	-5.62	5.27	0.17
66	Dy II	2	-6.94	3.94	0.02

(up to five times the solar value). While phosphorus is overabundant in the atmosphere of HD 100357, no gallium lines were detected.

4.5 Iron peak elements: Scandium to Nickel

With the exception of nickel, which possesses an abundance close to the solar value, the remaining six iron-peak elements exhibit significant overabundances, ranging from 1.0 to 2.0 dex.

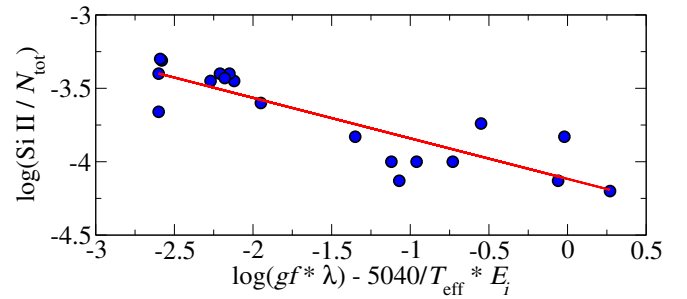


Figure 8. The scattered points indicate the individual Si abundances as a function of line strength in the atmosphere of HD 100357. The solid line represents a linear fit to the scattered data points, illustrating the declining abundance trend of Si with line strength, indicative of abundance stratification.

4.6 Neutron-capture elements

The abundance of Sr was measured using the Sr II $\lambda 4215.5 \text{ \AA}$ line and shows a significant enhancement of 3.67 dex compared to the solar value. Efforts to analyse other neutron capture elements, such as Y and Zr, were unsuccessful due to a lack of suitable spectral lines for accurate abundance determination.

4.7 Rare Earth Elements (REE)

HD 100357 displays significant overabundances of rare earth elements, including Nd, Gd, and Dy. Such high overabundances are a distinctive feature of upper main-sequence magnetic chemically peculiar stars. A similar chemical pattern was observed in the slowly rotating, strongly magnetic Ap star HD 144897, as reported by Ryabchikova et al. (2006). Owing to the high rotational velocity of HD 100357, we were able to reliably measure only 3 REE.

5 DOPPLER IMAGING

DI (Kochukhov 2017) is a robust technique to ascertain the topology of chemical spots on Ap stars. The technique involves mapping the surface inhomogeneities of a star given the observed line profile variations. For this, one needs high S/N ($S/N \sim 300$) line profiles of different chemical elements over adequate rotational phases. In the absence of sufficient S/N, we can use the LSD technique to produce a mean line profile with enhanced S/N appropriate for DI, where the S/N is increased proportional to the square root of the number of effective absorption lines.

For our study, we computed the LSD profiles for individual elements across the various observed phases for HD 100357. Out of the various sets of LSD profiles calculated for HD 100357, those derived for Cr and Fe exhibit the highest S/N, enabling us to map the surface spots of these elements using DI (Kochukhov 2016). To this end, we employed the DI code InversLSD (Kochukhov et al. 2014) to map the distributions of Cr and Fe in terms of the local line strength of the LSD profile. Following the approach of Fréour et al. (2023) and Semenko et al. (2024), the local spectra were computed under the assumption of a Milne-Eddington atmosphere, adopting a Voigt absorption profile. Limb darkening was treated according to the square-root law (Díaz-Cordoves & Gimenez 1992). The radial velocity offset and the projected rotational velocity were refined during the DI analysis by identifying the values that yielded the best fit to the observations.

The results of the DI calculations are presented in Fig. 9 and Fig. 10, which show the fits to the Cr and Fe LSD profiles and the corresponding surface line strength maps, respectively. This analysis yields $v \sin i$ values of 59.1 km s^{-1} and 59.8 km s^{-1} for Cr and Fe, respectively, consistent with other determinations in this study. The maps shown in Fig. 10 indicate that the two elements have qualitatively similar distributions, with Cr exhibiting a higher contrast surface structure. For both elements, the most prominent feature is a ring of relative underabundance encircling the star.

Fig. 11 shows the same Cr and Fe maps as in Fig. 10, but now in a rectangular format with the horizontal axis following the increasing rotational phase. Below the map is the phased *TESS* light curve. It is evident that the areas of relative underabundance for both elements coincide with the light minima. The result agrees with qualitative expectations from theory: a higher abundance increases UV absorption, leading to flux redistribution from UV to optical and NIR (Krtićka et al. 2015). Therefore, in the *TESS* bandpass, the star is brighter (fainter) where the element abundance is higher (lower).

6 DISCUSSION AND CONCLUSIONS

We report the detailed study of the Ap star HD 100357, which exhibits an abundance pattern consistent with the typical trends observed in other hot Ap stars. The star displays slight deficiencies in several light elements. The Fe-peak elements are overabundant by about one dex, except for Cr, Mn, and Co, those are overabundant by approximately two dex. The neutron capture element Sr is significantly enhanced, with an abundance exceeding the solar value by more than 3 dex. Additionally, the rare earth elements are markedly overabundant, ranging from three to five dex above their solar values. Also, we found evidence of atomic diffusion through the abundance stratification of Si at different depths. The overabundance of weaker lines of Si indicates a dominance of gravitational settling over radiative levitation.

The Ap stars are distinguished by their excess of Fe-peak elements and rare earth elements. A notable characteristic of hot (T_{eff}

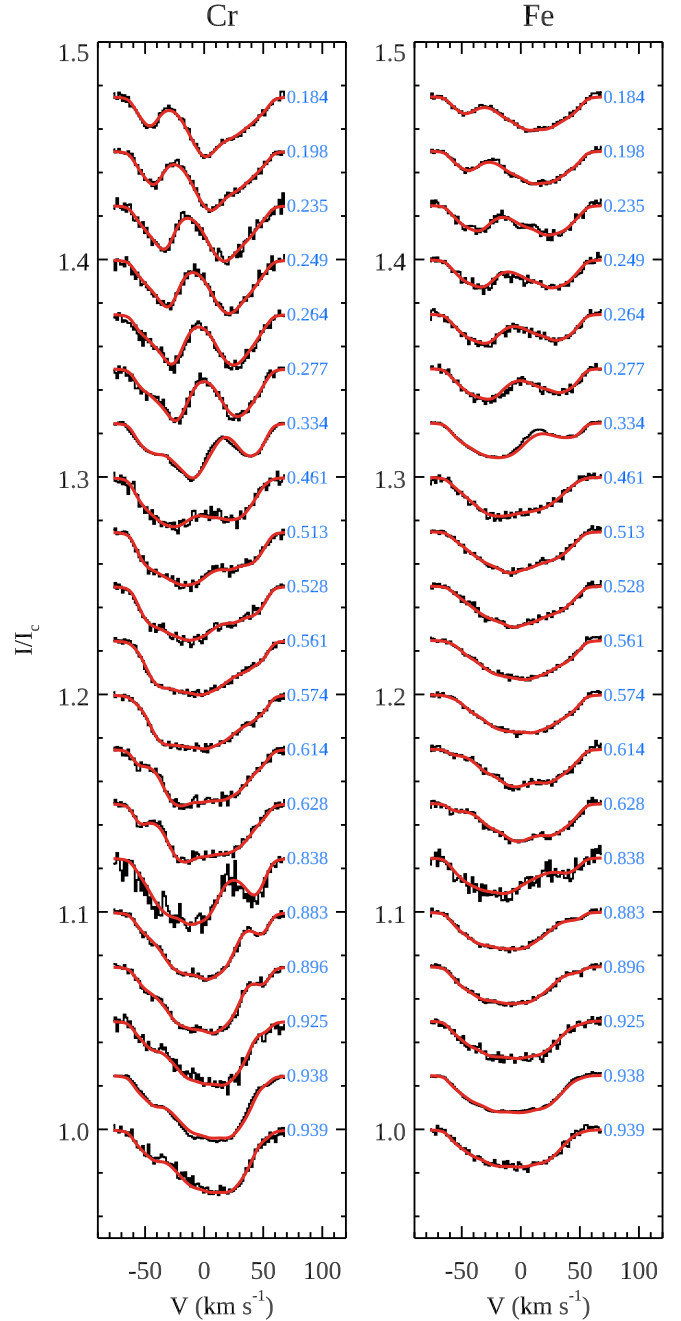


Figure 9. Comparison of the observed LSD profiles (histograms) and DI model LSD profiles (solid lines) of Cr and Fe. The profiles for different rotational phases are offset vertically. The phase values are shown to the right of the respective profiles.

> 10,000 K) counterparts in this class is an overabundance of Si. The deficiency of He is also seen in these stars, where the underabundance of He is correlated with the $\log g$ of the star (see Fig. 4 of Ghazaryan et al. 2018). Similarly, the Fe abundance is correlated to the T_{eff} of the star (see Fig. 5 of Ghazaryan et al. 2018). Thus, our target HD 100357 aligns well with all the characteristics of a hot magnetic CP star.

DI analysis revealed a ring of underabundance of Fe and Cr encircling the stars. Analogous to the DI results for other Ap stars (Kochukhov et al. 2019, 2022, 2023), these features presumably correspond to the equatorial region of the magnetic axis of a rather

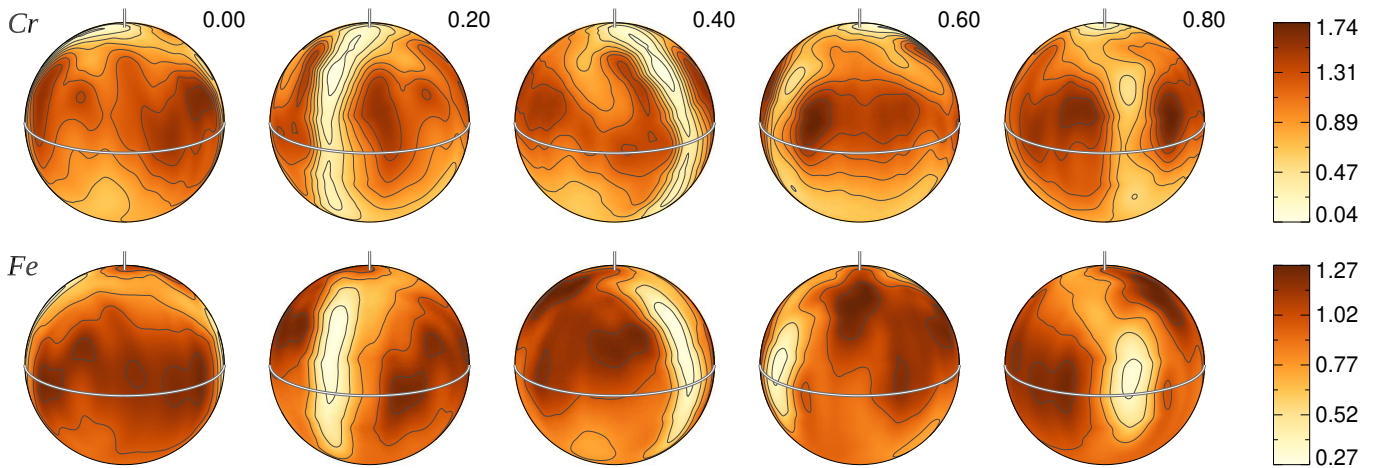


Figure 10. Surface distributions of Cr and Fe reconstructed with DI. The maps are given in terms of the logarithm of the local line strength. The star is shown at five rotational phases, indicated to the right of each column, at the inclination angle $i = 72^\circ$. In each spherical plot, the thick line shows the stellar equator, whereas the short bar corresponds to the rotational pole. We can notice a ring of underabundance regions inclined at 90° to the stellar equator.

weak dipolar-like magnetic field, which is inclined at about 90° to the stellar rotational axis.

Our quantitative examination of individual abundances and qualitative topological abundance distributions indicates that the star is likely to exhibit detectable magnetic fields. Consequently, HD 100357 can be classified as a hot Ap star exhibiting anomalies in Si, Sc, Ti, Cr, Mn, Fe, Co, and Sr, or CP2 as per Preston’s classification (Preston 1974). The study of these stars aids in comprehending the influence of magnetic fields and rotational forces on stellar atmospheres and chemical composition. Therefore, detailed spectropolarimetric observations are proposed to study the magnetic field topology of HD 100357.

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DATA AVAILABILITY

The data underlying this article will be shared upon request to the corresponding author. The *TESS* time-series flux data for the target star is publicly available from the NASA MAST archive (<https://mast.stsci.edu/portal/Mashup/Clients/Mast/Portal.html>). This work has made use of data from the European Space Agency (ESA) mission *Gaia* (<https://www.cosmos.esa.int/gaia>), processed by the Gaia Data Processing and Analysis Consortium (DPAC, <https://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for DPAC has been provided by national institutions, in particular those participating in the *Gaia* Multilateral Agreement. This research has also made use of the SIMBAD database, operated at CDS, Strasbourg, France. This research has used data, tools, or materials developed as part of the EXPLORE project, which has received funding from the European Union’s Horizon 2020 research and innovation programme under grant agreement No. 101004214.

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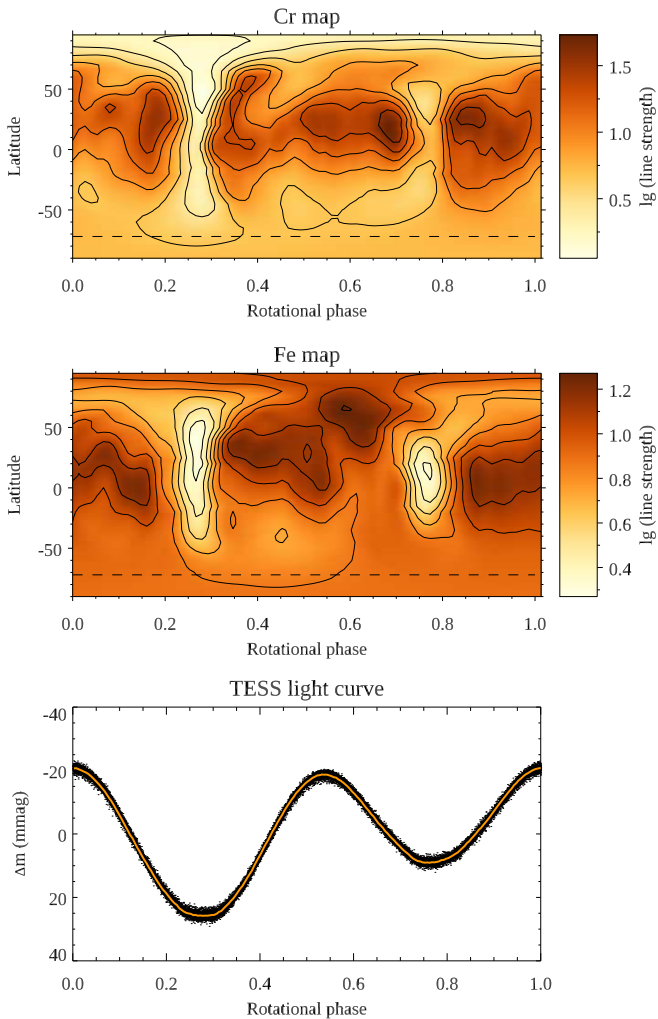


Figure 11. The top two panels show the surface distribution of Cr and Fe, respectively, as a 2D plane, where the x-axis gives the rotational phase and the y-axis represent the latitude. The dashed line represents the -72° latitude below which the surface of the star is not visible to the line of sight. At the bottom panel, the phased TESS light curve of HD 100357 is plotted for comparison.

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APPENDIX A: INDIVIDUAL LINE ABUNDANCES**APPENDIX B: CONTAMINATION**

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.

Table A1. The LTE abundances of HD 100357 derived from individual spectral lines of chemical species. The column $[X/H]$ represents the abundance $\log(A)_X$ relative to the solar abundances.

Atomic Number	Element	Wavelength (Å)	$\log(A)_X$	$[X/H]$	solar abundance
2	He II	5875	-3.1	-2	-1.1
2	He II	4471	-3.1	-2	-1.1
6	C II	4267.259	-4.7	-1.28	-3.42
8	O I	7771-7775	-4.27	-1.03	-3.24
12	Mg II	4481	-5.2	-0.7	-4.5
12	Mg II	7877.054	-5	-0.5	-4.5
12	Mg II	7896.366	-4.9	-0.4	-4.5
13	Al II	4663.046	-5.94	-0.35	-5.59
13	Al II	5593.2998	-6	-0.41	-5.59
13	Al II	7042.083	-6	-0.41	-5.59
14	Si II	4130.872	-3.83	0.65	-4.48
14	Si II	4130.894	-3.83	0.65	-4.48
14	Si II	4621.418	-3.45	1.03	-4.48
14	Si II	4621.7222	-3.45	1.03	-4.48
14	Si II	5041.0239	-3.74	0.74	-4.48
14	Si II	5055.9839	-4.13	0.35	-4.48
14	Si II	5056.3169	-4.13	0.35	-4.48
14	Si II	5688.817	-3.4	1.08	-4.48
14	Si II	5800.454	-3.31	1.17	-4.48
14	Si II	5806.731	-3.66	0.82	-4.48
14	Si II	5867.48	-3.4	1.08	-4.48
14	Si II	5868.4438	-3.4	1.08	-4.48
14	Si II	5957.5591	-4	0.48	-4.48
14	Si II	6371.3711	-4.2	0.28	-4.48
14	Si II	6660.532	-3.43	1.05	-4.48
14	Si II	6671.841	-3.6	0.88	-4.48
14	Si II	6699.431	-3.3	1.18	-4.48
14	Si II	7848.816	-4	0.48	-4.48
14	Si II	7849.722	-4	0.48	-4.48
14	Si III	4567.84	-3.5	0.98	-4.48
15	P II	4602	-5.6	0.98	-6.58
15	P II	6165.598	-5.5	1.08	-6.58
16	S II	5453.855	-4.3	0.58	-4.88
16	S II	5606.151	-4.2	0.68	-4.88
21	Sc II	4374.457	-7.15	1.7	-8.85
21	Sc II	4400.3892	-7.29	1.56	-8.85
21	Sc II	5526.79	-7.36	1.49	-8.85
22	Ti II	4163.644	-6.4	0.63	-7.03
22	Ti II	4443.801	-6.2	0.83	-7.03
22	Ti II	4468.493	-6.4	0.63	-7.03
22	Ti II	4571.972	-6.12	0.91	-7.03
24	Cr II	4145.7808	-4	2.38	-6.38
24	Cr II	4242.366	-4	2.38	-6.38
24	Cr II	4256.1079	-4.23	2.15	-6.38
24	Cr II	4275.5669	-4.1	2.28	-6.38
24	Cr II	4511.7749	-4.01	2.37	-6.38
24	Cr II	4539.595	-4.15	2.23	-6.38
24	Cr II	4592.052	-4.3	2.08	-6.38
24	Cr II	4616.6289	-4.02	2.36	-6.38
24	Cr II	4812.3369	-4.25	2.13	-6.38
24	Cr II	4824.127	-4.04	2.34	-6.38
24	Cr II	5420.925	-4.06	2.32	-6.38
25	Mn I	4030.7529	-4.45	2.1	-6.55
25	Mn II	4251.7168	-4.57	1.98	-6.55
25	Mn II	4252.9629	-4.57	1.98	-6.55
25	Mn II	4292.2368	-4.43	2.12	-6.55
25	Mn II	4478.6372	-4.43	2.12	-6.55
25	Mn II	4518.956	-4.53	2.02	-6.55

Table A1. Continued.

Atomic Number	Element	Wavelength (Å)	$\log(A)_X$	$[X/H]$	solar abundance
25	Mn II	4727.84	-4.6	1.95	-6.55
25	Mn II	4738.29	-4.5	2.05	-6.55
25	Mn II	4749.112	-4.51	2.04	-6.55
25	Mn II	4755.727	-4.31	2.24	-6.55
25	Mn II	4764.728	-4.45	2.1	-6.55
25	Mn II	4784.625	-4.5	2.05	-6.55
26	Fe II	4461.706	-3.25	1.23	-4.48
26	Fe II	4520.218	-3.52	0.96	-4.48
26	Fe II	4522.628	-3.27	1.21	-4.48
26	Fe II	4596.01	-3.34	1.14	-4.48
26	Fe II	4638.041	-3.17	1.31	-4.48
26	Fe II	4977.03	-3.11	1.37	-4.48
26	Fe II	5325.552	-3.26	1.22	-4.48
26	Fe II	5387.063	-3.46	1.02	-4.48
26	Fe II	5487.618	-3.46	1.02	-4.48
26	Fe II	5830.344	-3.39	1.09	-4.48
26	Fe II	5838.99	-3.3	1.18	-4.48
26	Fe II	5871.77	-3.26	1.22	-4.48
26	Fe II	6060.967	-3.3	1.18	-4.48
26	Fe II	7506.543	-3.46	1.02	-4.48
27	Co II	4569.25	-5	2.04	-7.04
27	Co II	4660.656	-5	2.04	-7.04
28	Ni I	4980.166	-5.8	-0.080	-5.72
28	Ni I	5080.527	-5.79	-0.070	-5.72
28	Ni I	5081.1069	-5.81	-0.090	-5.72
38	Sr II	4215.519	-5.41	3.67	-9.08
60	Nd III	4927.4877	-7.34	3.21	-10.55
60	Nd III	5294.1133	-7.2	3.35	-10.55
64	Gd II	4514.504	-5.74	5.15	-10.89
64	Gd II	4193.145	-5.5	5.39	-10.89
66	Dy II	4573.855	-6.92	3.95	-10.87
66	Dy II	4073.12	-6.95	3.92	-10.87

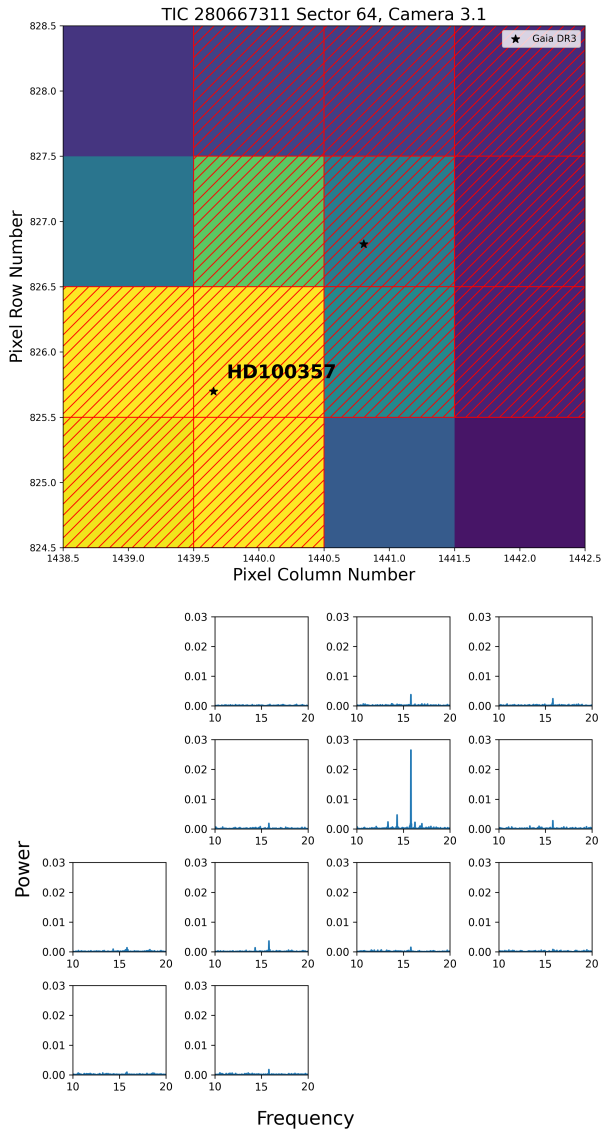


Figure B1. Top panel: The TESS pixels (4×4 , Sector 10) in the field of HD 100357 marked with a star symbol and labelled. To the top right, a nearby star is seen contaminating HD 100357. The hatched pixels show the masks used to plot the periodograms in the bottom panel. Bottom Panel: The corresponding periodograms of the selected masks. We can see the increased amplitude of the contaminating frequency at the pixel with the contaminating star and the amplitude drops towards the target star.

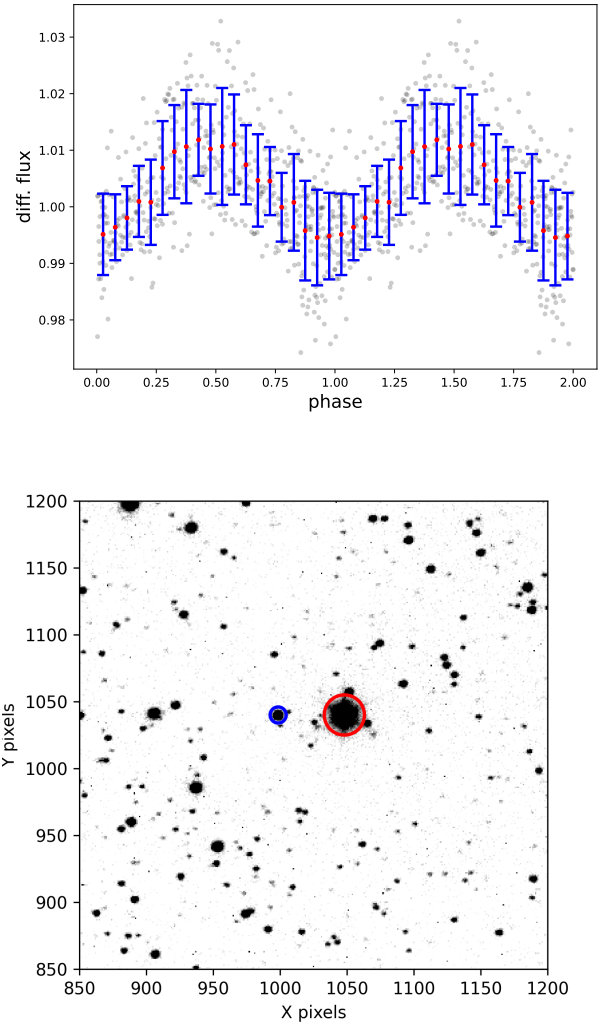


Figure B2. Top panel: The ground-based phased light curve of the contaminating star observed with TRAPPIST-South folded with the contaminating period. The red dots show the binned light curve, while the blue error bars are standard deviations in each bin. The grey dots in the background are the original phased light curve. Bottom: A cut-out of R-band CCD image in the field of HD 100357 taken with TRAPPIST-South. The larger red circle to the right is HD 100357 and the smaller blue circle to the left is Gaia DR3 5236626819707864192, which contaminates HD 100357 in TESS pixels.